

Reference Antenna Techniques for Canceling RFI due to Moving Sources

D. A. Mitchell,^{1,2} J. G. Robertson¹

We investigate characteristics of radio frequency interference (RFI) signals that can affect the excision potential of some interference mitigation algorithms. The techniques considered are those that modify signals from auxiliary reference antennas to model and cancel interference from an astronomical observation. These techniques can be applied in the time domain, where the RFI voltage is modeled and subtracted from the astronomy signal path (adaptive noise canceling), or they can be applied to the auto- and cross-correlated voltage spectra in the frequency domain (post-correlation canceling). For ideal receivers and a single, statistically stationary interfering signal, both pre- and post-correlation filters can result in complete cancellation of the interference from the observation. The post-correlation method has the advantage of being applied on 10s or 100s of millisecond time scales rather than 10s or 100s of nanosecond time scales. However this can be a disadvantage if the RFI transmitter location is changing, since the cross-correlated power measurements which link the interference power in the astronomy and reference signal paths can decorrelate. If the decorrelation is not too severe it can be allowed for, at the expense of a noise increase. The time domain adaptive cancelers are allowed to slightly vary their internal coefficients and adapt to changing phases during the integrations, which means that they avoid the decorrelation problem. However, the freedom to adapt also results in a noise increase. In this paper the ability of both types of cancelers to excise interference originating from a moving source is compared. The cancelers perform well on both observed and simulated data, giving complete cancellation.

1. Introduction

No matter what the astronomical application, radio frequency interference (RFI) is becoming an increasing problem in radio astronomy, and many methods for removing or suppressing the RFI are being proposed, evaluated and implemented. In most cases, the astronomer's targets are the correlations of signals from one or more antennas, and it is only these longer-term time averages of power that are wanted – there is no requirement to preserve voltage modulation. These applications are generally either finding the auto-correlation of signals from a single antenna (to measure the power spectrum of the as-

tronomy signal), or the cross-correlation of signals from more than one antenna (to measure the spatial coherence – or visibilities – of the astronomy signal). For details and specific examples see *Thompson et al.* [1986] and various chapters of *Taylor et al.* [1999].

Suppose that a sampled voltage stream consists of an additive mixture of components that are uncorrelated with each other. Suppose also that each component is statistically stationary so that if the component happens to be present in more than one voltage stream the phase difference and the ratio of the sampled amplitudes measured at two receivers are constant. Then as described in *Briggs et al.* [2000], if one of these correlated components is undesired and hindering our ability to probe cosmic components, it is possible to cancel this RFI from the power spectra, after the voltages have been correlated. Canceling RFI from correlations, referred to as “post-correlation canceling”, can offer many advantages over canceling the additive RFI voltage directly, particularly in regard to computational efficiency, since the canceling is performed on each correlation, 10s or 100s of times a second, rather than

¹School of Physics, University of Sydney, Bldg A28, Sydney, NSW 2006, Australia.

²CSIRO Australia Telescope National Facility, PO Box 76, Epping, NSW 1710, Australia.

each voltage sample, 10s or 100s of times a millisecond. Furthermore, it is possible to implement the technique in some current arrays with no modification, albeit at reduced array performance, for example *Kesteven* [2002]. Before considering cancellation techniques further, the signal itself needs to be briefly described.

Consider the voltage sequence sampled at an antenna as a combination of three complex additive components: receiver noise, $N(\nu, t)$; a noise-like cosmic component, $S(\nu, t)$; and interference, $I(\nu, t)$. Since each quasi-monochromatic spectral channel of the signal will be considered independently, the frequency labeling will be dropped to condense equations. After experiencing a phase shift, $\phi_m(t)$, due to the geometric delay of the signal relative to an arbitrary reference point, and being amplified and possibly phase shifted by a gain term, $G(t)$, the signal at the output of antenna m 's sampler at time step i , which is only measuring the voltage in the narrow spectral band centered at frequency ν , is

$$V_m(i) = S(i) + G_m(i)I(i)e^{j\phi_m(i)} + N_m(i), \quad (1)$$

where it is assumed that the signal has been amplified and delayed so that the cosmic signal is in phase with equal power at all of the receivers.

In the absence of the RFI component, one could detect and measure the amount of cosmic power by comparing the voltage sequences from two antennas, V_l and V_m , since the background noise is different for the two receivers. However, the presence of the interfering signal obscures the cosmic detection. If one were to compare each main antenna voltage sequence with voltage sequences from antennas that do not measure the cosmic signal (so only comparing the RFI components of various signals), then one could attempt to model and remove the RFI component in the comparison of V_l and V_m and recover the astronomy signal. These RFI references would be of the form

$$V_r(i) = G_r(i)I(i)e^{j\phi_r(i)} + N_r(i). \quad (2)$$

For any pair of antennas, the signals are compared by correlating the two voltage sequences together, that is, multiplying one signal by the complex conjugate of the other and accumulating the product for some accumulation time (~ 1 second). Uncorrelated components will multiply to give zero-mean noise which will average away as the number of sam-

ples accumulated approaches infinity, while a component that is present in both signals will correlate constructively, with an amplitude proportional to its power.

The correlated power terms in the obscured main antenna cross-correlation (so ignoring zero-mean noise terms) are

$$\begin{aligned} P_{lm} &= \langle V_l(i)V_m^*(i) \rangle \\ &= \sigma_S^2 + G_l G_m^* e^{j\phi_{lm}} \sigma_I^2 + \langle N_l(i)N_m^*(i) \rangle, \end{aligned} \quad (3)$$

where $\phi_{lm} = \phi_l - \phi_m$, σ_S^2 and σ_I^2 are the variances of the cosmic and interfering signals respectively, the asterisk superscript indicates a complex conjugation, the angular brackets represent the expectation operator which is approximated by a time average, and it has been assumed that the gain and phase terms are constant over the time interval. The cross-correlated receiver noise term, which should be zero-mean, has been included to remain general, since the following techniques also apply to auto-correlations where the noise is correlated against itself, that is, when $l = m$.

Post-correlation cancelers, which are briefly reviewed in the following section, estimate and then subtract the RFI power, $G_l G_m^* e^{j\phi_{lm}} \sigma_I^2$, from P_{lm} .

2. Post-Correlation Cancelers

The post-correlation technique described in *Briggs et al.* [2000] involves creating a model of the RFI in the main astronomy correlations using signals from a set of auxiliary reference antennas. Keeping in mind that each spectral channel is processed separately, the RFI model is a complex number with an amplitude and phase equal to the RFI component in (3).

Following *Briggs et al.* [2000], the post-correlation canceler estimates $G_l G_m^* e^{j\phi_{lm}} \sigma_I^2$ in P_{lm} using the closure relations, resulting in an RFI power model

$$\begin{aligned} M_{lm} &= \frac{P_{lr_1} P_{mr_2}^*}{P_{r_1 r_2}^*} \\ &= \frac{\langle G_l G_{r_1}^* \sigma_I^2 e^{j\phi_{lr_1}} \rangle \langle G_m G_{r_2}^* \sigma_I^2 e^{j\phi_{mr_2}} \rangle^*}{\langle N_{r_1} N_{r_2}^* \rangle^* + \langle G_{r_1} G_{r_2}^* \sigma_I^2 e^{j\phi_{r_1 r_2}} \rangle^*} \\ &\approx G_l G_m^* e^{j\phi_{lm}} \sigma_I^2, \end{aligned} \quad (4)$$

which is equal to the RFI term in P_{lm} .

In practice the expectation operators are not infinite in extent, and there is zero-mean noise centered on the correlated power in (4). So when M_{lm} is subtracted from P_{lm} the residual power will be

zero-mean noise, not zero, and there will always be an increase in noise over the situation where there was no RFI to begin with. This noise will average toward zero, and essentially result in an increase in system temperature (a decrease in the sensitivity of P_{lm}).

The closure relations which suggest the equality of the RFI in P_{lm} and the model M_{lm} only hold when the phase difference of the RFI signal remains constant for all antenna pairs. In practice though the geometric delays of the RFI signal between various antennas are not constant, due to the apparent motion of the RFI transmitter with respect to the array, which results in decorrelation of the cross-power measurements in (3) and (4), see for example *Thompson et al.* [1986]. The relative antenna gain to the RFI for each signal will also vary, but this is not considered here. If the only broken assumption is that of variable RFI geometric delays, and the delays are essentially constant over the time average, canceling before or after correlation will give very similar results. However, when the delays are changing appreciably during the integration, the different time scales on which the pre- and post-correlation algorithms are applied lead to differences in the canceling. To investigate the differences, we will first look at the effect of variable delays on the post-correlation algorithm, then pre-correlation cancelers which can track the changing delays during the time integrations.

3. Decorrelated Cross-Power

If the RFI signal geometric delays used in (3) or (4) are changing while the post-correlation RFI model is being calculated, the RFI power will be smeared over a range of delays, and the RFI model will be incorrect. This process, occurring because of the apparent motion of the array as it tracks cosmic sources - or the motion of the RFI source itself - is known as time-average smearing and the resulting loss of correlated power is known fringe rotation decorrelation, since the source is moving through fringes. If the gain terms stay approximately constant over the correlation averaging time, τ_A , then the correlated RFI power for arbitrary signals V_j and V_k , written now as a function of delay rather than phase, is

$$P_{I,jk} = \frac{G_j G_k^*}{\tau_A} \int_{t_0 - \frac{\tau_A}{2}}^{t_0 + \frac{\tau_A}{2}} e^{j2\pi\nu\tau_{jk}(t)} dt. \quad (5)$$

Let the geometric delay for the source at time t_0 be τ_{jk_0} seconds, and define τ'_{jk} to be the delay offset, $\tau_{jk} - \tau_{jk_0}$. Then if the rate of change of τ_{jk} with time is approximately constant over the integration and equal to $\Delta\tau_{jk}/\tau_A$ seconds of delay per second, the chain rule can be used to write the measured RFI power in (5) as a function of the change in geometric delay over the course of the integration

$$P_{I,jk}(\Delta\tau_{jk}) = \frac{G_j G_k^*}{\Delta\tau_{jk}} \int_{-\frac{\Delta\tau_{jk}}{2}}^{+\frac{\Delta\tau_{jk}}{2}} e^{j2\pi\nu\tau'_{jk}} d\tau'_{jk}. \quad (6)$$

The integral is the Fourier transform of a rectangular function, $\Pi(\tau'_{jk}/\Delta\tau_{jk})$, which has the value of unity inside the time interval and zero elsewhere, $\Pi(x) = \begin{cases} 1; & |x| \leq 1/2 \\ 0; & |x| > 1/2 \end{cases}$, which gives

$$P_{I,jk}(\Delta\tau_{jk}) = \text{sinc}(\nu\Delta\tau_{jk})P_{I,jk}(0). \quad (7)$$

The faster the delay is changing (and the higher the frequency), the more the cross-power estimate will be smeared out. The proportion of correlator output remaining after decorrelation is the ratio of $P_{I,jk}(\Delta\tau_{jk})$ to $P_{I,jk}(0)$, and for baseline jk is denoted F_{jk} . For the constantly varying geometric delay of $\Delta\tau_{jk}$ seconds per integration, we have

$$F_{jk} = \text{sinc}(\nu_0\Delta\tau_{jk}). \quad (8)$$

The effect of decorrelation can be seen in Figure 1. This data is Global Positioning System (GPS) satellite interference collected at two Australia Telescope Compact Array antennas separated by 4.4 km. The RF voltages were centered at 1575 MHz in a 4 MHz wide band, sampled with 4-bits, and summarized in *Bell et al.* [2001]. Figure 1a is a plot of cross-correlated power (scaled to be 1 for no decorrelation) for various integration lengths, which fits well to the theoretical curve given by (8) as the satellite moved across the sky.

So for RFI with an apparent motion relative to the sky reference frame, the post-correlation model for signals V_l and V_m given in (4) becomes

$$M_{lm} = \frac{F_{lr_1} F_{mr_2}}{F_{r_1 r_2}} G_l G_m^* \sigma_I^2, \quad (9)$$

while the RFI power in the main cross-power measurement will have decorrelated to $F_{lm} G_l G_m^* \sigma_I^2$, leaving residual power of

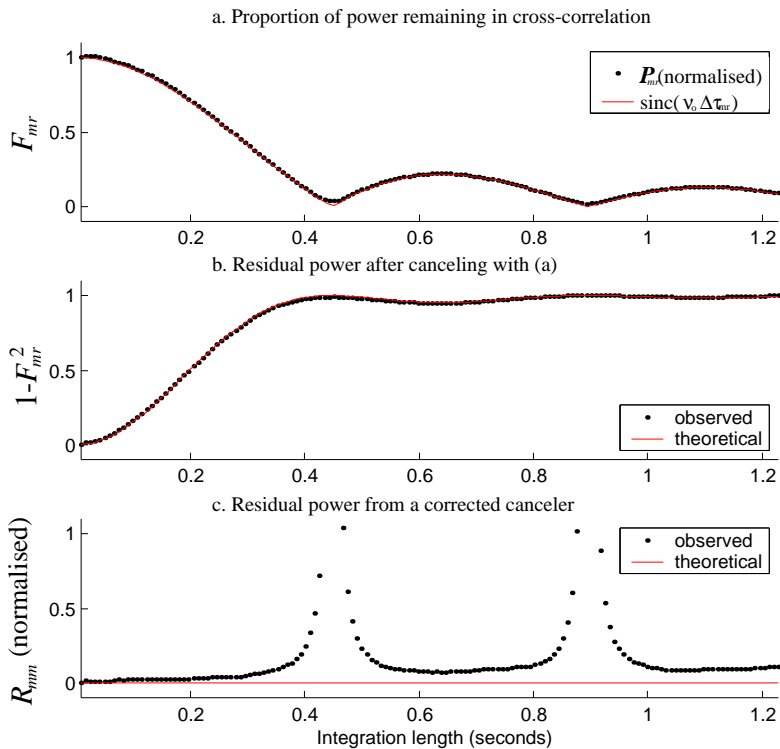


Figure 1. **a.** Proportion, F_{mr} , of the GPS power remaining after integration as a function of integration time τ_A . Dots show the measured power and the line shows the theoretical estimate. **b.** The proportion of residual RFI remaining in the output power after cancellation with M_{mm} . **c.** Proportion of residual RFI remaining in the output power after cancellation using the modified model, M'_{mm} .

$$R_{lm} = \left(F_{lm} - \frac{F_{lr_1} F_{mr_2}}{F_{r_1 r_2}} \right) G_l G_m^* \sigma_I^2. \quad (10)$$

Figure 1b shows the proportion of GPS satellite power remaining in the power spectrum of one ATCA antenna ($l = m$) after cancellation using both polarizations from the other ATCA antenna as the two references (one would typically not choose to place reference antennas 4.4 km from the main antennas, this is a worst-case example). Since the references are collocated, the only F_{jk} terms that will not be equal to 1 in (10) are F_{mr_1} and F_{mr_2} (written F_{mr} since they are equal). It is clear the cancellation becomes ineffective when decorrelation affects the cross-power measurement.

From (9) and (10) it can be seen that one can modify the post-correlation algorithm so that it takes the decorrelation into account

$$M'_{lm} = \frac{F_{lm} F_{r_1 r_2}}{F_{r_1} F_{mr_2}} M_{lm}, \quad (11)$$

where the prime on the model indicates that fringe rotation has been corrected for. This is shown in Figure 1c. Apart from the two regions where F_{mr} goes to zero and the noise is being infinitely magnified, the canceler achieves complete cancellation of the RFI, albeit with a small magnification of the noise. In those regions where most of the RFI power has been decorrelated away, so that one or more of the F_{jk} terms is close to zero, there is little that can be done to reverse the decorrelation. If it happens in the cross-power spectrum case that most of the RFI power has been decorrelated out of P_{lm} , then it might be better to do nothing as the RFI is canceling itself.

Figure 2a shows the residual spectra at three integration lengths when decorrelated RFI models, M_{mm} , were used, and Figure 2b shows the same spec-

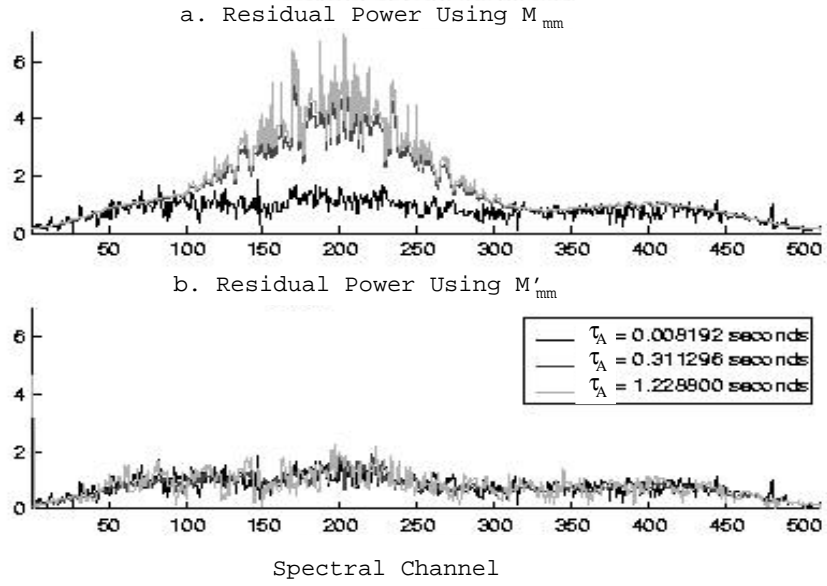


Figure 2. Residual power spectra after three integrations of length 0.008, 0.311, and 1.229 seconds, for **a.** the standard post-correlation filter, M_{mm} , and **b.** the modified post-correlation filter M'_{mm} .

tra when modified models, M'_{mm} , were used. The RFI is the peak centered at spectral channel 200, and the improvement achieved by scaling the weights is clear.

We will now show how, with constant geometric delays, equivalent canceling can be performed on the RFI voltages directly. This type of canceling needs to be applied much more often, however as discussed afterwards, it has the ability to track the changing delays.

4. Voltage Cancelers

Single reference pre-correlation adaptive noise cancelers are discussed in *Widrow and Stearns* [1985] and an early application to radio astronomy is described in *Barnbaum and Bradley* [1998]. A pre-correlation canceler which uses two reference signals to achieve equivalent cancellation to the post-correlation canceler is shown schematically in Figure 3. These dual-reference filters were suggested by *Briggs et al.* [2000]

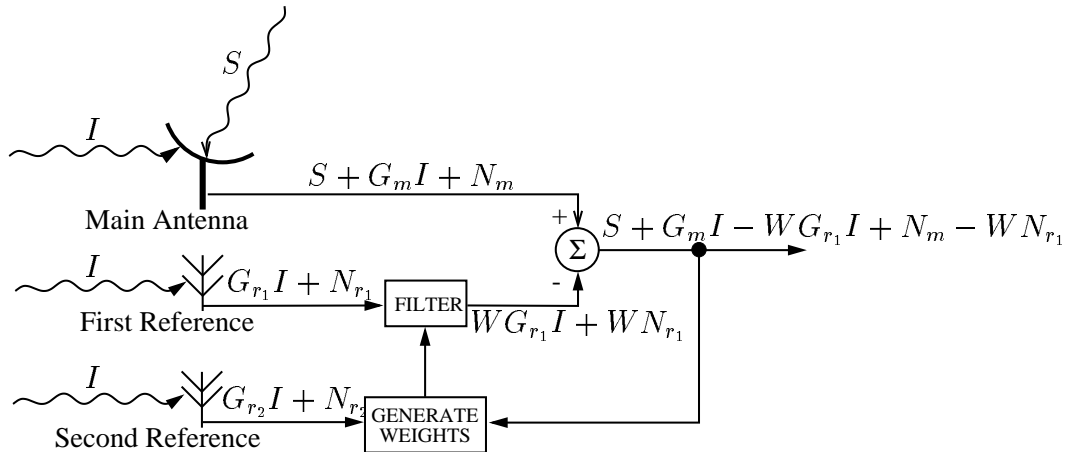


Figure 3. An adaptive canceler. The complex weights are those which set the cross-correlation between the output and the second reference spectrum to zero.

and are discussed in *Mitchell et al.* [2002]. Each spectral component of one reference signal is amplified and phase shifted by a weight so that its RFI component matches the RFI component in the main signal. A second reference is used to determine the required weight, which is chosen to set the correlation between the reference signal and the canceler output to zero. The second reference has receiver noise that is uncorrelated with the first, so that RFI is the only signal that could be in both the second reference voltage and the output voltage. Thus zeroing the cross-correlation means zeroing the RFI in the output. From Figure 3, the weight which sets the correlated power of the canceler output and reference r_2 to zero, $\langle (V_m - W_{mr_2} V_{r_1}) V_{r_2}^* \rangle = 0$, is

$$W_{mr_2} = \frac{\langle V_m V_{r_2}^* \rangle}{\langle V_{r_1} V_{r_2}^* \rangle}, \quad (12)$$

which, as in the post-correlation case, will be incorrect if the expectation values are estimated from cross-power measurements that have decorrelated due to changing RFI delays. The weights, which are optimal at any given instant, can be tracked rather than estimated statistically, as described in *Widrow and Stearns* [1985] and discussed in section 5. For the moment assume that we know the expectation

values in (12). Even though a RFI voltage model can be made to completely cancel the RFI in V_m , there will still be power in the output due to reference receiver noise (it is this noise which prompted the use of the second reference). To insure that the receiver noise added during the filtering of V_m will not be correlated with the receiver noise added during the filtering of V_l , we need to weight a different reference signal when main signal V_l is processed. Since the added receiver noise is all from the first reference receiver, swapping the references around for V_l means that noise from the second reference receiver will be added, as shown in Figure 4. The weight for V_l is

$$W_{lr_1} = \frac{\langle V_l V_{r_1}^* \rangle}{\langle V_{r_2} V_{r_1}^* \rangle}. \quad (13)$$

When W_{mr_2} and W_{lr_1} are used to filter V_m and V_l respectively, the output cross-power measurement is

$$P_{lm} = \langle (V_m - W_{mr_2} V_{r_1})(V_l - W_{lr_1} V_{r_2})^* \rangle. \quad (14)$$

As in the post-correlation case, this has completely removed the RFI but added zero-mean noise which decreases the sensitivity of the cross-power measurement. However, the canceler does not leave a corre-

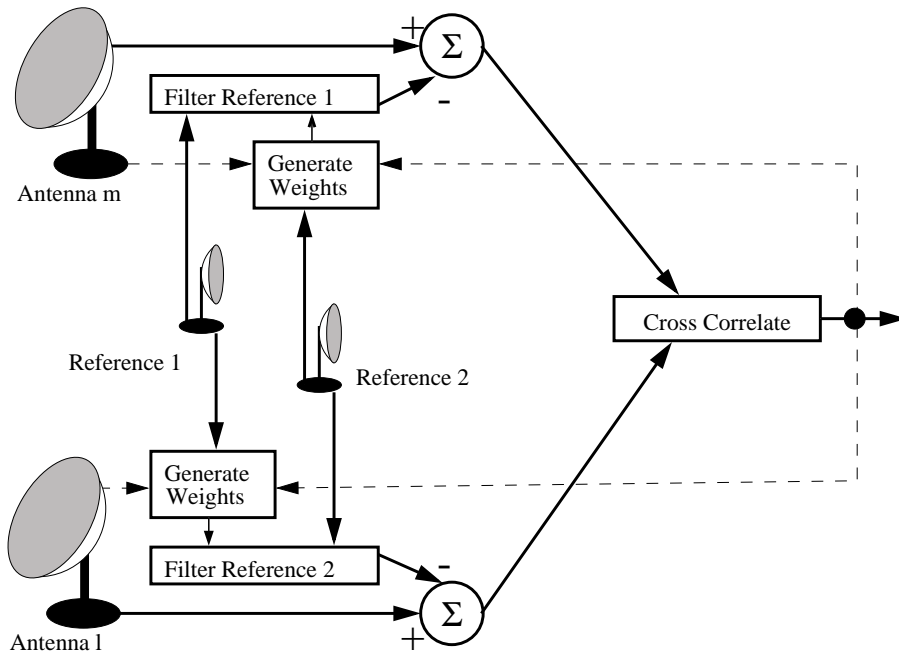


Figure 4. A double canceler setup where the model for each main signal is generated from a unique reference.

lated signal which would set a maximum achievable sensitivity level.

While the post-correlation algorithm described earlier can only account for the decorrelation after it has occurred, we now look at how the pre-correlation algorithm can be made to track any changing delays (one could of course account for the decorrelation in the pre-correlation weights by altering (12) and (13), with a result similar to that of post-correlation cancelers).

5. Adapting Time Domain Cancelers

When the optimal pre-correlation filter weights given by the expectation values in (12) and (13) are changing in time due to varying geometric delays, adaptive cancelers can attempt to keep up with the changing delays, by slightly modifying their weights at a rate equal to twice the bandwidth for Nyquist sampling, *Widrow and Stearns* [1985].

While the equations for RFI models given in section 4 give a single complex number for each spectral channel, the pre-correlation cancelers were actually applied in the time domain, directly to the voltages. In this situation, a set of reference voltage samples, delayed in time in both the positive and negative direction, are amplified by a real valued gain factor (the model weights), and then the weighted reference samples are added together to form the RFI model for the current sample. It should be pointed out that the time and frequency domain cancelers are essentially equivalent.

An example weight vector (refer to the simulation in the next section) is shown in Figure 5. This RFI is wide-band noise spread over the entire spectrum with equal power, so a suitable weight vector would be single peak at an appropriate delay (the negative of the geometric delay between the main and reference antennas) with an amplitude equal to the value the reference signal needs to be multiplied by to match the RFI in the main signal. This is indeed the weight vector that the algorithm gave – except for a small zero-mean noise floor (this “misadjustment” noise, due to the fact that the weights are being adapted using noisy power measurements, is discussed at length in *Widrow and Stearns*, 1985).

Given a set of starting weights and the associated output power, *Widrow and Stearns* [1985] give an estimate for the correct weights using the local power

gradient with respect to weights. For time step $i + 1$, the offset of the weight from the optimal value for the k^{th} delayed reference sample can be found from time step i

$$w_{i+1}(k) = w_i(k) - \frac{1}{2\sigma_r^2} \Delta_i(k), \quad (15)$$

where σ_r^2 is the variance of the reference signal and $\Delta_i(k)$ is the gradient of the canceler output power with respect to the k^{th} weight at the previous time step. Instead of attempting to find the optimal weights in one step (which would require using statistics as in section 4), the one step procedure can be replaced by an iterative process in which the weights are only slightly modified at each iteration, $w_{i+1}(k) = w_i(k) - \mu \Delta_i(k)$, where μ is a user defined constant called the adaptation gain factor which controls the rate of adaptation. *Widrow and Stearns* [1985] derive the equation for the power gradient term and show the weight adaptation to be equal to

$$w_{i+1}(k) = w_i(k) + 2\mu v_{out,m}(i)v_r(i-k), \quad (16)$$

where $v_{out,m}$ is the voltage at the output of the canceler for antenna m , and μ must be between 0 and $\frac{1}{(L+1)\sigma_r^2}$ for convergence to hold, where $L + 1$ is the number of delayed reference voltage samples used. Equation 16 shows that the iterative alteration of the weights really does form the correlation between the reference and the output and adapts until this correlation is zero. The value of μ controls the length of the integration, smaller values will smooth (reduce) the noise fluctuations in the weights, but increase the

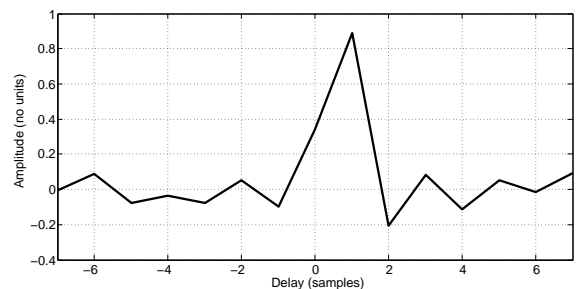


Figure 5. Example weight vector. The reference vector is delayed by each of the delays along the horizontal axis, multiplied by the corresponding amplitude, then all of the weighted copies are added together to give the model voltage estimate.

time constant for their response to changes (such as in the geometric delay).

Adaptive time domain cancelers can follow a moving source, avoiding decorrelation, but will eventually suffer a noise penalty when they have to adapt very quickly.

6. Simulated Comparisons

In this section the ideas discussed up to this point will be investigated using simulated wide-band Gaussian noise. To keep things simple and minimize the number of correlations affected by the changing delays (as in the GPS data from section 3), consider the situation of a single antenna observing a cosmic source in the presence of wide-band RFI. The quantity we wish to measure is the cosmic source power, which can be achieved by calculating the auto-correlation of the received voltage sequence, but this will also contain receiver noise and RFI power. Power measurements are made while implementing both pre- and post-correlation cancellation algorithms, both of which will be affected by the changing delays, as discussed in the previous sections. Since an auto-correlation is being considered, the signal index, l , becomes m in all of the previous equations and the RFI in the main signal auto-correlation will suffer no decorrelation, $F_{mm} = 1$. Also, consider two reference signals coming from orthogonal polarizations of a single reference antenna (antenna r), so that the references are collocated and the RFI in their cross-correlation will also suffer no decorrelation. So long as the RFI is 100% polarized (and not orthogonal to the polarization of any of the receivers), then $F_{rr} = 1$, and the two cross-power measurements between signal m and each of the reference signals are decorrelated down to a fraction F_{mr} of the incident power. Equations (9), (10) and (11) reduce to

$$\begin{aligned} M_{mm} &= F_{mr}^2 G_m G_m^* \sigma_I^2, \\ R_{mm} &= (1 - F_{mr}^2) G_m G_m^* \sigma_I^2, \text{ and} \\ M'_{mm} &= \frac{1}{F_{mr}^2} M_{mm} = P_{I,mm}. \end{aligned} \quad (17)$$

To investigate these equations, three voltages signals were generated, V_m , V_{r1} and V_{r2} , each of which was broadband Gaussian noise. The main voltage

was the sum of a broadband Gaussian RFI sequence and a unique Gaussian noise sequence simulating receiver noise. The reference signals contained a delayed and amplified version of the RFI and their own unique receiver noise sequences. As the voltages were generated the delay between the RFI in the main voltage sequence and the RFI in the reference voltage sequences was slightly incremented, simulating a moving source.

Figure 6 shows RFI power in the cross-correlation of signal m and one of the references, along with the RFI power after the decorrelation has been corrected for. The decorrelated spectrum, P_{mr} in Figure 6a, is divided by F_{mr} (the curved line), to give P'_{mr} , shown in Figure 6b. The power level in each frequency channel has been correctly reset to be centered on the expected (decorrelation free) value. The noise in the power measurement has increased, which is particularly clear in the higher frequency channels where more amplification was needed.

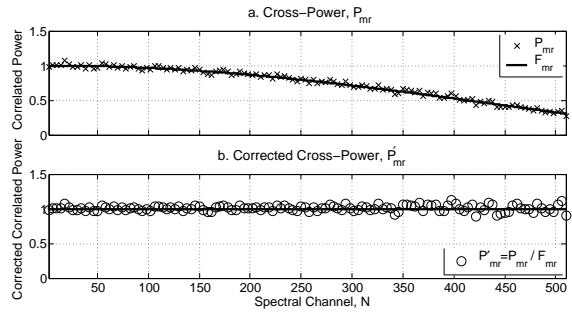


Figure 6. a. Normalized cross-correlated RFI power, P_{mr} , on the baseline between the main and reference antennas which has been decorrelated; b. The cross-correlated power after allowing for decorrelation, P'_{mr} .

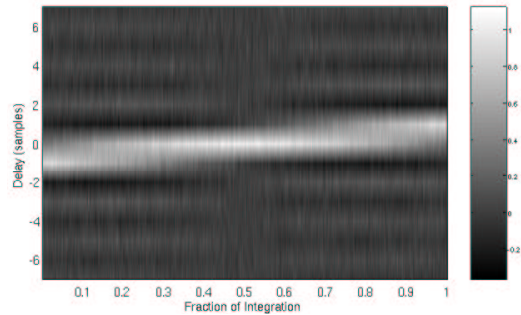


Figure 7. Weight vectors as a function of integration number as the integration progresses. Grey-scale indicates the amplitude of the weights.

Figure 7 is a gray-scale map of the adaptive pre-correlation weight vector (refer to Figure 5) displayed vertically as the integration progresses through the voltage samples. In the simulations $L = 16$, that is, 8 negative and 8 positive delays. The adaptation gain factor was set to be 1 percent of the delay change over the integration (1% of 1.5 Nyquist samples), multiplied by the maximum value allowed; $\mu = \frac{1.5}{100} \frac{1}{(L+1)\sigma_r^2}$. The geometric delay has changed by 1.5 sample lengths over the course of the integration. The gray value of each pixel indicates the amplitude of the weight. As the delay changes, the main peak in the weight vector follows it. As the fringe rate increases, the adaptive filter can easily keep up (as long as the stability requirements are still met).

The residual power after canceling is shown in Figure 8. The dashed line shows the theoretical level of reference receiver noise added during filtering when the RFI transmitter is stationary and the geometric delay is constant; zero-mean noise which will continue to average toward zero as the integration length is increased. The adaptive filter results (indicated by the dots), follow this theoretical limit across the entire spectrum, indicating that the filter is indeed tracking the changing delays. The solid line and crosses represent the post-correlation canceler which

has not been modified to account for decorrelation. The excess power above the dotted line is residual interference which will not average down any further. The residual power after post-correlation canceling using M'_{mm} , which has accounted for fringe rotation (Figure 6b), is indicated by open circles and continues to average down toward the theoretical limit as the frequency channel and decorrelation increase. The residual power in the adaptive and corrected post-correlation filtered spectra does not contain RFI, and will average down with the dashed light blue line, albeit with a larger noise floor when more fringe rotation correction is needed (that is, the higher frequency channels in Figures 6 and 6).

So while the post-correlation canceler cannot be made to track an interfering signal's geometric delay faster than the integration length, the resulting decorrelation of the cross-correlations in the filter equations can often be accounted for.

7. Conclusions

Both the pre- and post-correlation RFI cancelers discussed in this paper are adversely affected by interfering sources that are moving relative to cosmic sources. This is due to changing geometric de-

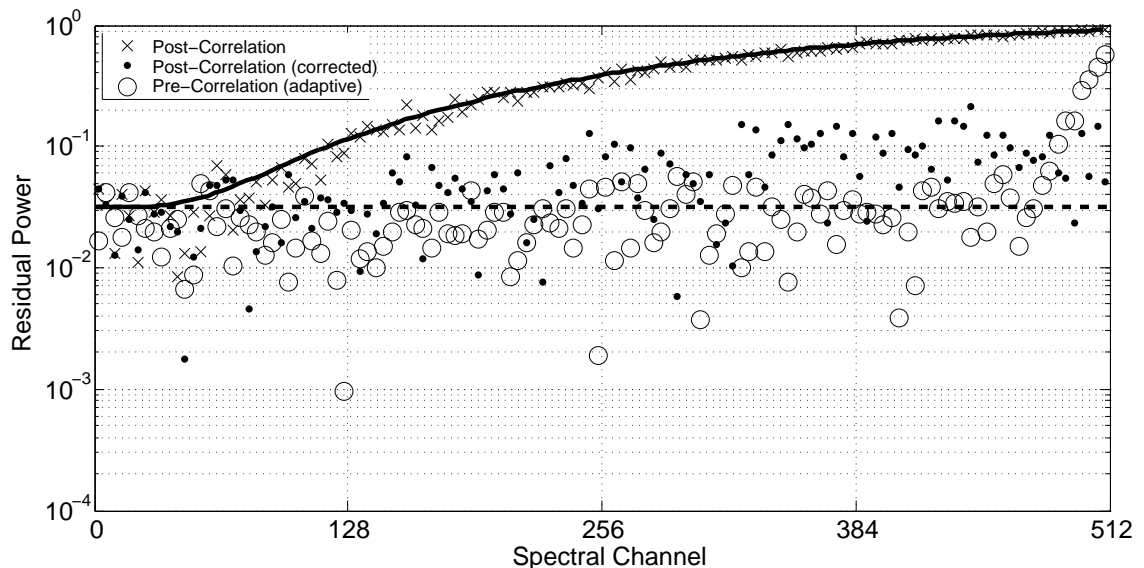


Figure 8. Residual power after filtering a wide-band RFI signal (normalized to unity for no filtering). The lines indicate the theoretical amount of residual power remaining; the solid line for a moving source and the dashed line for a stationary source. Points indicate the simulation results, dots for the adaptive pre-correlation canceler, crosses and open circles for post-correlation canceler with and without decorrelation corrections respectively.

lays causing decorrelation of the cross-power measurements used to generate the RFI models. If the rate of change of the geometric delays is known and the amount of decorrelation is only slight, the models can be corrected by amplifying the cross-power measurements back to their theoretical levels. This process can be avoided in pre-correlation cancelers, since they can track the correct models much faster (closer the voltage sample rate which is to the order of the signal bandwidth). Both of these filters result in extra noise relative to stationary RFI situation.

One should keep in mind the particular case of terrestrial interference from a fixed source. In this situation the interferometer is fringe tracking the celestial sphere so the RFI fringe rate is the negative of the sidereal fringe rate. Since the correlator is inserting this phase term, one always knows the fringe rate for each pair of antennas, and for antenna pair jk , F_{jk} can calculate exactly. The required level of suppression and added complexity, however, may dictate that shorter integration times are necessary to eliminate the need for the decorrelation corrections.

Acknowledgments. We acknowledge Mike Kesteven, Frank Briggs, Bob Sault and Lawrence Cram for many enlightening discussions on the topic, and the Scientific Committee on Frequency Allocations for Radio Astronomy and Space Science (IUCAF) and the University of Sydney, for funding to attend the Workshop in Mitigation of Radio Frequency Interference in Radio Astronomy.

References

- Barnbaum, C., and R. F. Bradley (1998), A New Approach to Interference Excision in Radio Astronomy: Real-Time Adaptive Cancellation, *Astronomical Journal*, 116, 2598–2614.
- Bell, J. F., P. J. Hall, W. E. Wilson, R. J. Sault, R. J. Smegal, M. R. Smith, W. van Stratten, M. J. Kesteven, R. H. Ferris, F. H. Briggs, G. J. Carrad, M. W. Sinclair, R. G. Gough, J. M. Sarkissian, J. D. Bunton, and M. Bailes (2001), Base Band Data for Testing Interference Mitigation Algorithms, *Publications of the Astronomical Society of Australia*, 18, 105–113.
- Briggs, F. H., J. F. Bell, and M. J. Kesteven (2000), Removing Radio Interference from Contaminated Astronomical Spectra Using an Independent Reference Signal and Closure Relations, *Astronomical Journal*, 120, 3351–3361.
- Kesteven, M. J. (2002), Adaptive Filters for RFI Mitigation in Radioastronomy, New Technologies in VLBI, IVS Symposium in Korea, Gyeong-Ju, Korea.
- Mitchell, D. A., M. J. Kesteven, and R. J. Sault (2002), Post Correlation Versus Real Time Adaptive RFI Cancellation, XXVIIth General Assembly of the International Union of Radio Science, Maastricht, The Netherlands.
- Taylor, G. B., R. A. Perley, and C. L. Carilli, editors (1999), *Synthesis Imaging in Radio Astronomy II*, 180, Astronomical Society of the Pacific Conference Series, Astronomical Society of the Pacific, San Francisco, CA.
- Thompson, A. R., J. M. Moran, and G. W. Swenson (1986), *Interferometry and Synthesis in Radio Astronomy*, Wiley-Interscience, New York, NY.
- Widrow, B., and S. D. Stearns (1985), *Adaptive Signal Processing*, Prentice Hall, Englewood Cliffs, NJ.
- D. A. Mitchell, School of Physics, University of Sydney, Bldg A28, Sydney, NSW 2006, Australia. (mitch@physics.usyd.edu.au)
- J. G. Robertson, School of Physics, University of Sydney, Bldg A28, Sydney, NSW 2006, Australia. (jgr@physics.usyd.edu.au)

(Received _____.)